

A Charge Membrane Paradigm

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the work in progress
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Introduction

- Emparan, Suzuki, Tanabe (EST) and collaborators have recently noted that the classical dynamics of black holes simplifies when D , the dimensionality of space time is taken to be large .
- The spacetime becomes effectively flat at distance of the order $\delta r = r_0/D$ from the horizon of a D dimensional RN black hole in this limit, where $r_0 =$ radius of its horizon.
- the RN black holes in large D are characterised by **two** length scales: r_0 and δr , and this separation also shows up in the spectrum of their quasinormal modes.

Introduction

- EST showed that most of the quasinormal modes are 'heavy', with frequencies $\mathcal{O}(1/\delta r)$ while the remaining are anomalously light, with frequencies $\mathcal{O}(1/r_0)$.
- For all these modes the imaginary parts are of the same order in D as their real parts.
- That means in a dynamical process the heavy modes decay much more quickly than their light counterparts. And the dynamics at late times $t \gg \delta r$ is governed by a nonlinear interacting theory of the light modes which are just a handful in number.

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What is this membrane anyway?

- As mentioned earlier, the black hole in large D has non-flat metric only up to $\mathcal{O}(r_0/D)$ outside the horizon, while the region inside the horizon is inconsequential since it is causally disconnected from the region of interest.
- Therefore these spacetimes are shown to be in one-to-one correspondence with the configurations of a non-gravitational codimension one surface propagating in flat D dimensional space. This surface is called a **membrane**.
- The dynamical degrees of freedom of this membrane are: the shape of the membrane, a velocity field u which lies entirely on the membrane and a scalar charge density Q .

and what are these 'equations'?

- It turns out that the metric and the gauge field within distance of $\mathcal{O}(\delta r)$ outside the horizon (i.e. the **membrane region**) can be written entirely out of these quantities (shape function, u and Q) and their derivatives.
- The derivatives of the shape function that will be useful here are n (the unit normal to the membrane), K (the extrinsic curvature of the membrane) and \mathcal{K} (Trace of K).

- These solutions are regular at the membrane (i.e. the horizon) upto the subleading order in $1/D$ if and only if the membrane satisfied the equations

$$\left(\frac{\nabla^2 u}{\mathcal{K}} - (1 - Q^2) \frac{\nabla \mathcal{K}}{\mathcal{K}} + u \cdot K - (1 + Q^2)(u \cdot \nabla)u \right) \cdot \mathcal{P} = 0,$$

$$\frac{\nabla^2 Q}{\mathcal{K}} - u \cdot \nabla Q + Q \left(\frac{u \cdot \nabla \mathcal{K}}{\mathcal{K}} - u \cdot K \cdot u \right) = 0,$$

where ∇ = the covariant derivative on the membrane world volume,
 and $\mathcal{P}_{AB} = \eta_{AB} - n_A n_B + u_A u_B$.

These are called the '**membrane equations of motion**'.

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Zeroth order ansatz

We know that the RN blackhole boosted with velocity u

$$ds^2 = ds_{flat}^2 + \left((1 + Q^2) \frac{1}{\rho^{D-3}} - Q^2 \frac{1}{\rho^{2(D-3)}} \right) O_M O_N dx^M dx^N,$$

$$A_M = \frac{\sqrt{2} Q O_M}{\rho^{D-3}},$$

$$O = n - u, \quad u = \text{const}, \quad u \cdot u = -1, \quad \rho = \frac{r}{r_0},$$

$$r^2 = P_{MN} x^M x^N, \quad P_{MN} = \eta_{MN} + u_M u_N, \quad n = r_0 d\rho, \quad u \cdot n = 0.$$

Solves the Einstein-Maxwell equations. So we perturb around this solution in $1/D$.

Metric and gauge field upto subleading order

$$h_{MN} = F(\rho)O_M O_N + H_{MN}^{(T)} + 2O_{(M}H_{N)}^{(V)} + H^{(S)}O_M O_N + H^{(Tr)}\mathcal{P}_{MN}$$

$$A_M = \sqrt{2}Q \rho^{-(D-3)} O_M + \left(A^{(S)}O_M + A_M^{(V)} \right),$$

where

$$F(\rho) = \left[(1 + Q^2)\rho^{-(D-3)} - Q^2\rho^{-2(D-3)} \right],$$

$$\mathcal{P}_{MN} = \eta_{MN} - O_M n_N - O_N n_M + O_M O_N,$$

$$\mathcal{P}^{MN} H_N^{(V)} = \mathcal{P}^{MN} A_N^{(V)} = 0, \quad \mathcal{P}^{MN} H_{MQ}^{(T)} = 0, \quad \mathcal{P}^{MN} H_{MN}^{(T)} = 0,$$

where

$$\begin{aligned}
 A_M^{(V)} &= - \left(\frac{2\sqrt{2}}{D} \right) Q \rho^{-D} \left[D(\rho - 1) (\mathfrak{Y}_{(1)} - Q^2 \mathfrak{Y}_{(2)}) \right. \\
 &\quad \left. - Q^2 [1 + \log(1 - \rho^{-D} Q^2)] \mathfrak{Y}_{(2)} \right]_M + \mathcal{O} \left(\frac{1}{D} \right)^2, \\
 A^{(S)} &= \left(\frac{1}{D} \right) \left[\sqrt{2} Q D(\rho - 1) \rho^{-D} \mathfrak{S}_{(1)} \right. \\
 &\quad \left. + 2\sqrt{2} \left(\frac{Q^3}{1 - Q^2} \right) \rho^{-D} \Upsilon_A(\rho) \mathfrak{S}_{(2)} \right] + \mathcal{O} \left(\frac{1}{D} \right)^2.
 \end{aligned}$$

$$H_{MN}^{(T)} = \left(\frac{2}{D}\right) \log(1 - Q^2 \rho^{-D}) \mathfrak{T}_{MN} + \mathcal{O}\left(\frac{1}{D}\right)^2,$$

$$H_M^{(V)} = \left\{ Q^2 \left[(F(\rho) - \rho^{-(D-3)}) + (F(\rho) - 1) \log(1 - Q^2 \rho^{-D}) \right] \mathfrak{B}_{(2)M} \right. \\ \left. - D(\rho - 1)F(\rho) [\mathfrak{B}_{(1)} - Q^2 \mathfrak{B}_{(2)}]_M \right\} \left(\frac{2}{D}\right) + \mathcal{O}\left(\frac{1}{D}\right)^2.$$

$$H^{(S)} = -\sqrt{2}Q \rho^{-D} A^{(S)} + \left(\frac{1}{D}\right) \left[\rho^{-(D-3)} - F(\rho) \right] \\ + \left(\frac{2}{D}\right) \rho^{-D} \left[Q^2 D(\rho - 1) \mathfrak{S}_{(1)} + \Upsilon_H(\rho) \mathfrak{S}_{(2)} \right] + \mathcal{O}\left(\frac{1}{D}\right)^2,$$

$$H^{(Tr)} = \mathcal{O}\left(\frac{1}{D}\right)^3,$$

where

$$F(\rho) = \left[(1 + Q^2)\rho^{-(D-3)} - Q^2\rho^{-2(D-3)} \right],$$

$$\Upsilon_A(\rho) = \int_0^{D(\rho-1)} dx \log(1 - Q^2 e^{-x}),$$

$$\begin{aligned} \Upsilon_H(\rho) = & \left[(\rho^D - Q^2) \log(1 - Q^2 \rho^{-D}) - (1 - Q^2) \log(1 - Q^2) \right. \\ & \left. + Q^2 \left(\frac{1 + Q^2}{1 - Q^2} \right) \Upsilon_A(\rho) \right]. \end{aligned}$$

Membrane equations as dynamical equations

- Total number of dynamical variables on the membrane are $D - 1$.
- Total number of membrane equations is also $D - 1$.
- Hence these equations define an initial value problem for these variables which can be solved for the dynamics of the membrane and hence for the corresponding black hole solution.
- It can be easily shown that the RN black hole is an exact solution of the membrane equations of motion.

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Light quasinormal modes of an RN black hole

- Now we perturb the membrane equations about the RN black hole solution.
- Solutions of the linearized equations have the frequencies

$$\omega_l^r = \frac{-i(l-1) \pm \sqrt{(l-1)(1-lQ_0^2)}}{1+Q_0^2},$$

$$\omega_l^Q = -il,$$

$$\omega_l^v = \frac{-i(l-2)}{1+Q_0^2},$$

- where Q_0 is the charge density parameter for the RN black hole and l is the spherical harmonic number .

- So, for given l there are only **3** scalar (two ω_l^r and one ω_l^Q) and **1** vector (ω_l^v) quasinormal modes.
- So for every l the light quasinormal modes are finite in number, as desired.
- In the chargeless limit ($Q_0 \rightarrow 0$) these frequencies reduce to the frequencies reported in **1406.1258** and **1504.06613** for LQNM of the Schwarzschild black hole (and ω_l^Q is absent).

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Defining conserved currents

- Our method gives the metric and gauge field inside the membrane region.
- While outside the membrane region the gauge field and the deviation of the metric from flatness are very small, so Einstein-Maxwell equations are linear in this region.
- So for the solutions in these two regions to be consistent, the membrane region solution must produce a conserved charge current and stress tensor which produce the gauge field and the metric resp. outside the membrane region.

Charge Current

- The constraint Maxwell equations $\nabla_\mu F^{\mu\rho} = 0$ suggest that there is a conserved current on the membrane

$$J_Q^\mu = \frac{\sqrt{g}}{\sqrt{g_{red}}} F^{\mu\rho}.$$

- This current is evaluated using the metric and gauge field listed above.

$$J = -\frac{2\mathcal{K}Q}{\rho^{D-2}}u + \frac{Q}{\rho^{D-3}}(u \cdot \nabla u) \cdot P + \frac{\mathcal{K}}{D\rho^{D-3}}dQ \\ - \frac{Q}{\rho^{D-3}}D(\rho - 1) \left((1 - Q^2) \frac{\nabla\mathcal{K}}{\mathcal{K}} + (1 + Q^2)u \cdot \nabla u \right) \cdot P$$

Membrane equations from the Charge current conservation

- The conservation equation for this current on the membrane yields

$$\begin{aligned} \nabla \cdot J = & \frac{1}{\rho^{D-3}} \left[\frac{\nabla^2 Q}{\mathcal{K}} - u \cdot \nabla Q - Q \left(\frac{u \cdot \nabla \mathcal{K}}{\mathcal{K}} - u \cdot K \cdot u \right) \right] \\ & - \frac{Q}{\rho^{D-3}} D(\rho - 1) \left[(1 - Q^2) \left(\frac{\nabla^2 \mathcal{K}}{\mathcal{K}^2} - \frac{u \cdot \nabla \mathcal{K}}{\mathcal{K}} \right) \right. \\ & \left. - (1 + Q^2) \left(\frac{u \cdot \nabla \mathcal{K}}{\mathcal{K}} - u \cdot K \cdot u \right) \right] \end{aligned}$$

- Note that the first line is the scalar membrane equation while the rest is proportional to the divergence of the vector equation.
- Hence the membrane equations are actually demanded by the conservation of charge current.

Stress tensor and entropy current

- Similarly we believe that the membrane equations indicate the conservation of Brown-York stress tensor.

$$T^{\mu\nu} = \frac{\mathcal{K}}{D} u^\mu u^\nu + \mathcal{O}\left(\frac{1}{D}\right)$$

- And the entropy current can be defined on the membrane by Hodge dualising the area form on it. (See [0803.2526](#))

$$J_S^\mu = u^\mu + \mathcal{O}\left(\frac{1}{D}\right)$$

- These currents can be used to compute the charges which can be verified to satisfy the first law of BH thermodynamics.

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Radiation

- The charge current and the stress tensor defined on the membrane can be considered as the sources for the gravitational and Maxwell fields outside the membrane.
- These sources convoluted with the respective retarded Green's functions will presumably give the gravitational and electromagnetic radiation.
- Although the sources are of unit magnitude, the radiation is very feeble, with the magnitude of order $1/D^D$. This is a kinematical fact about large number of dimensions, and has nothing to do with other specifics of the system.

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Major references:

- “A Charged Membrane Paradigm at Large D ”, 1511.03432
- Bhattacharyya et. al. 1504.06613 (It's free of charge.)
- EST: 1406.1258 etc. (Or just type 'large D ' in google)
- Entropy current: 0803.2526